

# VLT observations of the highly ionized nebula around Brey2<sup>★</sup>

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## Abstract.

We present the first high resolution He II  $\lambda 4686$  images of the high excitation nebula around the WR star Brey 2 in the LMC. This nebula presents a striking morphology: a small arc-like feature some 3.6 pc in radius is particularly prominent in the He II  $\lambda 4686$  line. We further discover a previously unknown faint He II emission that extends over an area of  $22 \times 17$  pc<sup>2</sup>. An even fainter He II emission is apparently associated with the interstellar bubble blown by the progenitor of Brey2. The total He II flux corresponds to an ionizing flux of  $4 \times 10^{47}$  photons s<sup>-1</sup>. H $\alpha$ , [O III], and He I  $\lambda 5876$  images and long-slit spectra are also examined in this letter, enabling us to investigate the detailed physical properties at various locations of the nebula.

**Key words.** Stars: individual: Brey2 - ISM: bubbles - ISM: abundances - ISM: supernova remnants

## 1. Introduction

Except for Planetary Nebulae, high excitation features, such as nebular He II emission, are not expected in H II regions. Indeed, normal O stars do not emit enough hard UV photons to produce detectable nebular He II emission. However, in the last two decades, seven objects of the Local Group were found to harbour such nebular He II emission (Garnett et al. 1991). Wolf-Rayet (WR) stars are the ionizing sources of five of them. These stars are hot, evolved massive stars that possess strong winds, but until then they were generally believed to be unable to excite such highly ionized nebulae. Nebular He II emission was also discovered farther away, in star-forming galaxies (e.g. I Zw 18, see French 1980) and WR stars were again thought to be responsible for it (Schaerer 1996). This type of emission thus seems intimately linked to WR stars in most cases, and the analysis of these peculiar highly ionized objects can help us put tighter constraints on the poorly known extreme UV ionizing fluxes of these stars (Crowther 2002). Only the He II nebulae of the Local Group are situated close enough to enable a detailed study.

We study here the nebula associated with Brey2 (or BAT99-2 in Breysacher et al. 1999), a WN2b(h) star in the Large Magellanic Cloud (Foellmi et al. 2003). This WR star has blown a small bubble whose expansion velocity reaches  $16 \text{ km s}^{-1}$  (Chu et al. 1999). Moderate chemical enrichment

was detected in this bubble (Garnett & Chu 1994). Using this result, Chu et al. (1999) suggested that the circumstellar bubble was probably currently merging with the interstellar bubble blown by the progenitor of the WR. Nebular He II emission near Brey2 was first detected by Pakull (1991). Imaging of the nebula in the He II lines is crucial in establishing the distribution of the emission in detail, in order to evaluate the total flux, and to better understand the origin of the ionization. To date, only low resolution, low signal/noise He II images exist (Melnick & Heydari-Malayeri 1991). Thanks to observations made with the Very Large Telescope (VLT), we present here the first high resolution He II  $\lambda 4686$  images and a detailed spectral analysis of this peculiar highly ionized nebula.

## 2. Observations

We obtained CCD images of Brey2 and its associated nebula with the FORS instrument installed on the 8 m VLT-UT3 in 2002 January. The images were taken through seven filters (H $\alpha$ , [O III], He I  $\lambda 5876$ , He II  $\lambda 4686$ , plus three continuum filters centered on 4850, 5300, and 6665 Å) for exposure times of  $3 \times 100$ s,  $3 \times 100$ s,  $9 \times 400$ s,  $9 \times 400$ s,  $9 \times 60$ s,  $9 \times 60$ s, and  $3 \times 100$ s, respectively. The seeing was  $\sim 1''$ . The data were reduced with IRAF using standard methods for overscan and bias subtraction and flatfielding. Stellar sources were removed using the photometric and astrometric information obtained from the continuum images (details are reported by Nazé et al., in preparation). The few remaining faint stars were either removed individually or not considered for flux determinations. Fig. 1 presents a three color image of the whole field; while Fig. 2 shows a close-up on Brey2 in the four nebular filters.

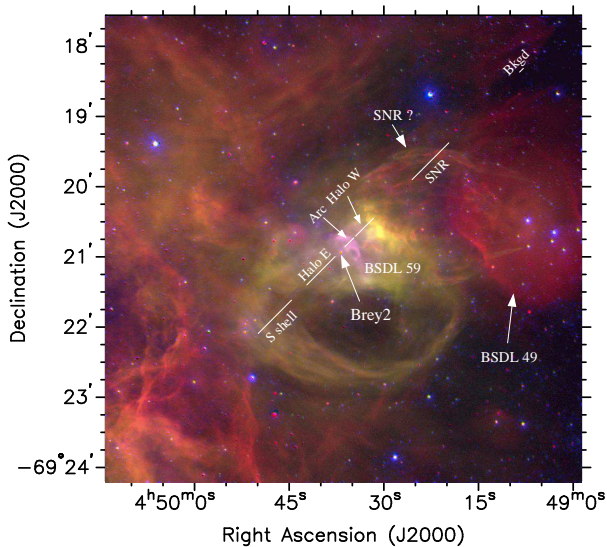
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<sup>★</sup> Based on observations collected at the European Southern Observatory, Cerro Paranal, Chile (ESO No. 68.C-0238(A,B)).

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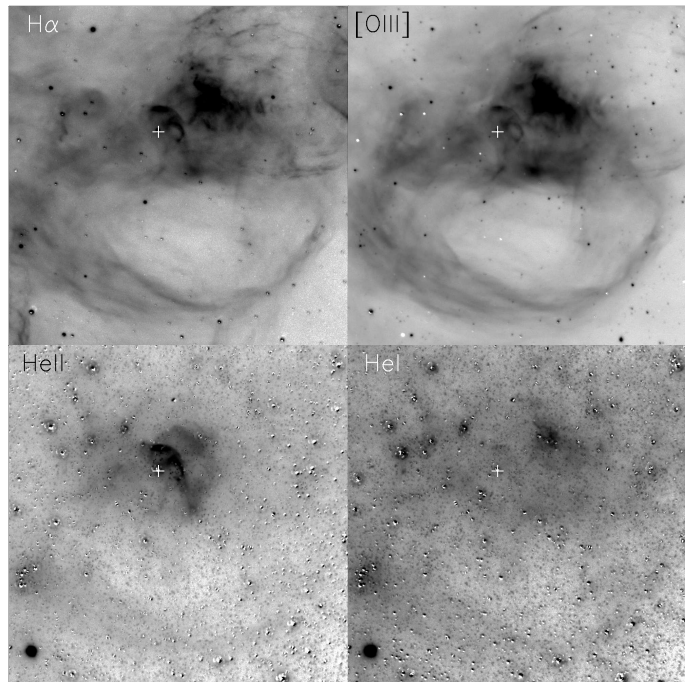
**Fig. 1.** Color image of the Brey2 field. Red, green and blue correspond to continuum subtracted  $H\alpha$ ,  $[O\text{ III}]$ , and  $\text{He II}$  images, respectively. The different regions used for spectral analysis are marked by a solid line. Features discussed in the text are labelled.

During the same observing run, we also obtained long-slit spectra of Brey2 with the same instrument. We used the 600B and 600V gratings to obtain a blue spectrum covering 3700–5600 Å ( $R \sim 800$ ) and a red spectrum covering 4500–6850 Å ( $R \sim 1000$ ), respectively. The  $1.3'' \times 6.8''$  slit was tilted by  $45^\circ$  and centered on Brey2. The spatial resolution was  $\sim 1.2''$  and the spectral resolution, as measured from the FWHM of the calibration lines,  $7\text{Å}$ . The spectra were reduced and calibrated in a standard way using IRAF. For flux calibration, we observed several standard stars from Oke (1990) and used the mean atmospheric extinction coefficients for CTIO reduced by 15%. Sky subtraction was done using a small region of the spectra where the nebular emission is the lowest. Only a few residuals remained for the brighter sky lines (e.g.  $[O\text{ I}] 5577\text{Å}$ ).

### 3. The high excitation nebula surrounding Brey2

#### 3.1. Morphology

As seen in Fig. 1, the nebula surrounding Brey2 is quite complex. The bubble blown by the WR star is visible to the NW of the star as a small, quite bright arc (and arrow)-like structure, of external radius  $\sim 15''$ . To the south of the star, a large shell,  $\sim 100''$  in radius, may in fact be the interstellar bubble blown by the progenitor of Brey2 (Garnett & Chu 1994). Though with a low S/N, the  $\text{He I}$  image appears well correlated with  $H\alpha$ , at least for the brightest features. But the  $[O\text{ III}]$  image seems less filamentary than the  $H\alpha$  image at the SNR's position (see below) and at the SE of the field-of-view. It also shows several additional features compared to the  $H\alpha$  image. All  $H\alpha$  features on the eastern side of the nebula possess a lower  $[O\text{ III}]/H\alpha$  ratio than the central nebulosities surrounding Brey2. BSDL 49 (Bica et al. 1999), an oval-shaped structure  $160''$  west of Brey2 and  $140'' \times 100''$  in size, presents an extremely low  $[O\text{ III}]/H\alpha$  ratio. These low ratios suggest that the regions are not ionized



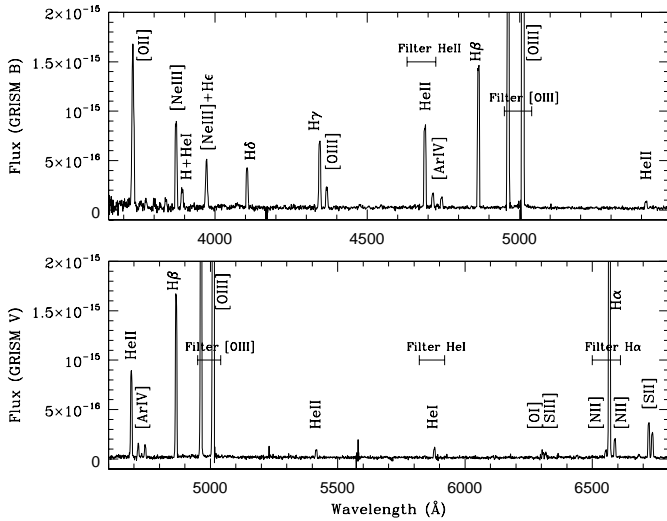
**Fig. 2.** FORS continuum subtracted  $H\alpha$ ,  $[O\text{ III}]$ ,  $\text{He II}$ , and  $\text{He I}$  images of Brey2 and its close surroundings. The images are  $200'' \times 200''$ . A white cross indicates the WR star's position.

by early O or WR stars but most probably by cooler stars, e.g. late O or early B stars.

In contrast, the  $\text{He II}$  image appears strikingly different (see Fig. 2). The arc-like structure emits strongly in  $\text{He II}$ , and a fainter halo  $\sim 90'' \times 70''$  in size surrounds the star. These  $\text{He II}$  emissions appear well centered on Brey2. The shell directly south of the star also emits  $\text{He II}$ , but rather faintly. This shell may represent the limit of the  $\text{He II}$  Strömgren sphere. Since the density south of Brey2 appears much lower than in the other directions, some  $\text{He II}$  ionizing photons could escape the direct surroundings of Brey2 and ionize this shell. Finally, note that the very faint halo seems to extend to the bright  $H\alpha$  regions visible at the lower left of Fig. 1. However, since we do not have spectra of this region, we can not confirm that the nebular emission actually comes from  $\text{He II}$ , and not from other nebular lines, e.g.  $[\text{Ar IV}]$ .

We used the calibrated long-slit spectra to get the ADU-flux conversion factor of the images. The total  $\text{He II}$  emission arc+halo is about  $5 \times 10^{-13} \text{ erg cm}^{-2} \text{ s}^{-1}$ , or a luminosity of  $3 \times 10^{35} \text{ erg s}^{-1}$  at the distance of the LMC (50 kpc), and after a reddening correction of  $A_V = 0.7 \text{ mag}$  (see §3.2). This corresponds to a flux of  $4 \times 10^{47} \text{ He}^+$  ionizing photons  $\text{s}^{-1}$ . Comparing to the models of Smith et al. (2002), this flux is compatible with a 90–100 kK WR star, consistent with the WN2 classification of Brey2.

To the NW of Brey2, a SNR was diagnosed by its high-velocity motions and confirmed through nonthermal radio emission and its high  $[\text{S II}]/H\alpha$  ratio (Chu et al. 1999). Between BSDL 49 and the features associated with Brey2, filamentary structures, characteristic of SNRs (Chen et al. 2000), prove the presence of this SNR candidate. Two rather bright  $H\alpha$  filaments



**Fig. 3.** FORS calibrated, sky-subtracted blue (upper) and red (lower) spectra of the arc near Brey2. The lines analysed in Table 1 are labelled, and a solid line shows the bandpass of the imaging filters. Note that the He II filter also includes the [Ar IV] lines, whilst the H $\alpha$  filter also contains the [N II] lines.

seem to delineate the SNR to the north and south. They form a region some  $\sim 50''$  in radius. Our long-slit spectra taken at that position show clear indications of an enhanced [S II]/H $\alpha$  ratio, typical of SNRs (see below), and a lower [O III]/H $\alpha$  ratio.

### 3.2. Spectrophotometry

The Balmer decrement H $\alpha$ /H $\beta$  was used to derive the interstellar extinction, assuming the theoretical case B decrement at  $T = 15000$  K and  $30000$  K (Storey & Hummer 1995). Some faint stars sometimes affected the nebular spectra. In such cases, we correct the measured emission line strengths by a 2 Å equivalent width for the Balmer absorptions (McCall et al. 1985) before estimating the reddening. To deredden the line ratios, we used the extinction law from Cardelli et al. (1989) with  $R_V = 3.1$ . Spectra were extracted from five regions of interest for detailed analysis (Fig. 1): the first one includes the arc of the arc-like structure (12'' long without the star); the second and third ones explore the eastern and western He II halos (34'' and 22'' long, respectively); the fourth examines the eastern part of the large southern shell (40'' long); and the fifth investigates the properties of the SNR candidate situated to the NW of the star (44'' long). The extinction-corrected line ratios for these regions are listed in Table 1.

Using the NEBULAR package under IRAF (Shaw & Dufour 1995), we then derived temperatures, densities and abundances in these regions (see Table 2). To get Helium abundances, we used the emissivities computed by Benjamin et al. (1999) for He I and by Storey & Hummer (1995) for H and He II. The presence of He II suggests that higher ionization states than O $^{2+}$  exist as well. This can actually explain the slightly smaller [O III]/H $\alpha$  ratio of the arc-like structure compared to its direct surroundings. Following Garnett & Chu (1994), we derived the total oxygen abundance using :

$$\frac{N(O)}{N(H)} = \frac{N(O^{0+}) + N(O^{+}) + N(O^{2+})}{N(H^{+})} \times \left( 1 + \frac{N(He^{2+})}{N(He^{+})} \right)$$

Compared to the LMC abundances (Russell & Dopita 1990), no sign of enrichment in helium can be detected for the He II nebula. But its oxygen abundance is slightly lower (max 35%), and its N/O ratio slightly higher (min 30%) than the LMC's trend. This is compatible with what might be expected for a nebula enriched by a stellar wind. Using a closer background for analysing the spectrum of the arc does not change this conclusion. On the other hand, the SNR shows rather low helium and oxygen abundances, but a normal N/O ratio.

## 4. Conclusion

We have presented the first high resolution images of the nebula around Brey2, one of the 7 He II nebulae of the Local Group. He II is particularly emitted by a small arc-like structure, some  $\sim 15''$  in radius, that surrounds the star. This arc corresponds to the bubble blown by the WR star. Fainter He II emission is present in a surrounding halo, and still fainter He II emission is detected in the large interstellar bubble blown by the progenitor of Brey2 towards the south of the star. We derived an He $^{+}$  ionizing flux of  $4 \times 10^{47}$  photons s $^{-1}$ , which appears consistent with the latest theoretical models for very early WN stars. We also confirm the results of Garnett & Chu (1994) regarding the composition of the nebula: with a larger N/O ratio and a slightly smaller O abundance, the small nebula associated with Brey2 shows only modest enrichment by stellar winds. We have also identified the SNR in the FORS images, and have analysed its spectrum. The bright - thus dense - H $\alpha$  region situated between Brey2 and the SNR may be the result of the compression of the ISM between these two expanding structures.

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**Table 1.** Dereddened line ratios with respect to  $H\beta=100.0$  (see text for details). The extinction was calculated using the  $H\alpha/H\beta$  ratio. Uncertainties, estimated from the signal/noise in the lines and the calibration errors, are given in parentheses.

Line	$A(\lambda)/A_V$	Arc	Halo E	Halo W	S shell	SNR
[O II] 3727	1.54	132 (14)	268 (28)	225 (23)	344 (36)	853 (89)
[Ne III] 3868	1.50	69 (7)	96 (10)	93 (9)	85 (8)	84 (8)
H8 + He I 3889	1.50	17 (2)	19 (2)	19 (5)	17 (2)	
[Ne III] + He $\epsilon$ 3967	1.47	38 (4)	42 (4)	44 (4)	31 (3)	23 (2)
H $\delta$	1.43	30 (3)	30 (3)	30 (3)	28 (3)	27 (3)
H $\gamma$	1.35	50 (4)	52 (4)	51 (4)	50 (4)	52 (4)
[O III] 4363	1.34	16 (1)	16 (1)	16 (1)	13 (1)	28 (2)
He I 4471	1.30			3.6 (0.3)		
He II 4686	1.22	56 (4)	21 (01)	24 (2)	8.1 (0.6)	
[Ar IV] 4711 <sup>a</sup>	1.21	8.2 (0.6)		3.8 (0.3)		
[Ar IV] 4740	1.20	6.9 (0.5)		3.8 (0.3)		
H $\beta$	1.16	100.	100.	100.	100.	100.
[O III] 4959	1.13	234 (17)	313 (22)	318 (23)	272 (20)	165 (12)
[O III] 5007	1.12	691 (49)	945 (67)	944 (67)	804 (58)	489 (35)
He II 5412	1.02	4.5 (0.4)		2.1 (0.2)		
He I 5876	0.93	5.2 (0.5)	10 (1)	9.6 (0.8)	11 (1)	9.1 (0.8)
[O I] 6300	0.86	3.1 (0.3)	4.9 (0.5)	7.3 (0.7)		25 (2)
[S III] 6312	0.86	2.6 (0.3)	3.7 (0.4)	3.0 (0.3)		
[O I] 6364	0.85			2.4 (0.2)		8.8 (0.8)
[N II] 6548	0.82	3.4 (0.3)	7.9 (0.8)	5.5 (0.6)	10 (1)	17 (2)
H $\alpha$ <sup>b</sup>	0.82	279 (28)	279 (28)	279 (28)	279 (28)	270 (27)
[N II] 6583	0.82	10 (1)	22 (2)	19 (2)	30 (3)	54 (5)
He I 6678	0.80	1.7 (0.2)	2.7 (0.3)	2.9 (0.3)	4.3 (0.5)	
[S II] 6716	0.79	16 (2)	31 (3)	29 (3)	40 (4)	93 (10)
[S II] 6731	0.79	12 (1)	22 (4)	21 (2)	29 (3)	68 (7)
$F(H\beta)$ (erg cm <sup>-2</sup> s <sup>-1</sup> )		1.16e-14	1.30e-14	1.93e-14	1.25e-14	7.47e-15
$A_V$ (mag)		0.65 (0.22)	0.70 (0.22)	0.72 (0.22)	0.62 (0.22)	1.06 (0.22)

<sup>a</sup> The small pollution due to He I 4713 was corrected using the strength of He I 5876 and the theoretical He I ratios from Benjamin et al. (1999).

<sup>b</sup> Pollution due to He II 6560 negligible.

**Table 2.** Derived physical properties of the 5 regions. The abundances were calculated assuming  $n_e = 100$  cm<sup>-3</sup>. If several lines of the same ion exist, the abundance shown here is an average of the abundances derived for each line.

Prop.	Arc	Halo E	Halo W	S shell	SNR
$T_e$ [O III] (K)	16100±750	14300±600	14100±650	14000±600	29800±2700
$n_e$ [S II] (cm <sup>-3</sup> )	<350	<310	< 310	<280	< 370
$n_e$ [Ar IV] (cm <sup>-3</sup> )	950-4000		3200-7000		
He <sup>+</sup> /H <sup>+</sup> × 10 <sup>2</sup>	4.44±0.30	7.78±0.53	7.70±0.40	10.3±0.7	7.17±0.62
He <sup>2+</sup> /H <sup>+</sup> × 10 <sup>2</sup>	4.91±0.26	1.80±0.13	2.19±0.12	0.70±0.05	
→ He/H × 10 <sup>2</sup>	9.35±0.40	9.58±0.55	9.89±0.42	11.0±0.7	7.17±0.62
O <sup>0+</sup> /H <sup>+</sup> × 10 <sup>6</sup>	1.26±0.12	2.80±0.27	4.47±0.31		2.53±0.17
O <sup>+</sup> /H <sup>+</sup> × 10 <sup>5</sup>	0.90±0.09	2.64±0.27	2.32±0.24	3.63±0.38	3.25±0.34 <sup>b</sup>
O <sup>2+</sup> /H <sup>+</sup> × 10 <sup>4</sup>	0.61±0.03	1.12±0.05	1.17±0.05	1.02±0.05	0.13±0.01
→ O/H × 10 <sup>4</sup>	1.51±0.09	1.73±0.08	1.86±0.08	1.48±0.06	0.48±0.01 <sup>b</sup>
N <sup>+</sup> /H <sup>+</sup> × 10 <sup>6</sup>	0.69±0.05	1.95±0.14	1.54±0.11	2.64±0.19	1.38±0.10
(→ N/O × 10 <sup>2</sup> ) <sup>a</sup>	7.66±0.98	7.39±0.93	6.64±0.84	7.29±0.92	4.24±0.53 <sup>b</sup>
S <sup>+</sup> /H <sup>+</sup> × 10 <sup>7</sup>	2.44±0.19	5.70±0.45	5.51±0.41	7.72±0.59	6.07±0.44
S <sup>2+</sup> /H <sup>+</sup> × 10 <sup>6</sup>	1.17±0.11	2.36±0.23	2.01±0.19		
Ar <sup>3+</sup> /H <sup>+</sup> × 10 <sup>7</sup>	4.31±0.23		2.93±0.16		
Ne <sup>2+</sup> /H <sup>+</sup> × 10 <sup>5</sup>	1.55±0.15	3.01±0.30	3.04±0.30	2.85±0.28	0.50±0.05

<sup>a</sup> Assuming  $N(N^+)/N(O^+) = N(N)/N(O)$

<sup>b</sup> To consider with caution : the calculation of the [O I] abundance was made at 20 000 K.